

# On the potentials yielding cosmological scaling solutions

Ana Nunes<sup>1,2\*</sup> & José P. Mimoso<sup>1,3†</sup>

<sup>1</sup>Dep. Física, Faculdade de Ciências, Ed C8, Campo Grande  
Universidade de Lisboa, 1700 Lisboa, Portugal

<sup>2</sup>CMAF and <sup>3</sup>CFNUL, Av. Prof. Gama Pinto, 2, 1699 Lisboa Codex, Portugal  
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## Abstract

In the present work we perform a phase-plane analysis of the complete dynamical system corresponding to a flat FRW cosmological models with a perfect fluid and a self-interacting scalar field and show that every positive and monotonous potential which is asymptotically exponential yields a scaling solution as a global attractor.

## I. INTRODUCTION

Scaling solutions of cosmological models with a self-interacting scalar field have attracted a lot of interest recently. It has been advocated that scalar field cosmologies offer a way to reconcile the standard cosmological model with the large scale structure (LSS) data [1] and the remarkable results of the supernovae observations [2] that indicate that the baryonic matter density parameter is bound to be  $\Omega_m \lesssim 0.3$  and that the universe is positively accelerated at present. Scalar fields would thus act like a decaying cosmological constant, but would otherwise have the advantage of evading some of the longstanding difficulties faced by the latter [3,4].

Rather than knowing the exact solutions of the scalar field cosmologies, what matters most is the asymptotic behaviour of models. The self-similar solutions that have been found when the potential of the self-interacting scalar field is exponential are particular remarkable in this sense [5–8]. They correspond to an equilibrium between the different kinds of energy that compete. Indeed the energy density of the scalar field scales with the same rate as that of the barotropic matter (some authors [9] refer to these solutions as *trackers* and denote *scaling* solutions those for which the energy density of the scalar field evolves as a power of that of matter; in what follows we shall adopt this terminology). Moreover these solutions attract all the other phase space trajectories in the case of flat space models and, hence, provide the late-time asymptotic behaviour for the scalar field cosmologies under consideration. This gives, on the one hand, a possible answer for why a non-vanishing

\*email: anunes@lmc.fc.ul.pt

†email:jpmimoso@alf1.cii.fc.ul.pt

scalar field does not introduce radical changes with respect to the usual Einstein-de Sitter rate of expansion of the universe. On other hand, it may contribute an explanation to the difference between the actual density of matter and the critical energy density of the spatially flat isotropic models. Furthermore, the scalar field component would also fulfill the convenient role of delaying the time of matter-radiation equality which would help fitting the power spectrum of large scale structure [3,19,4].

In the literature we find mainly two sorts of potentials underlying the tracking and the scaling behaviour: the exponential potentials and a class of power law potentials [10,9,11] respectively. However it is worth noticing that the solutions corresponding to the latter set of potentials were only shown to hold in the regime where the perfect fluid component fully dominates the expansion and the energy density of the scalar field is negligible. Regarding the solutions termed trackers [9]: the self-adjustment of the behaviour of the energy density of one or more scalar fields with that of matter has been investigated in Friedmann-Robertson-Walker (FRW) models [5–8,12,13] both with and without curvature, in spatially homogeneous, but anisotropic models [14], and in FRW models in scalar-tensor gravity theories (also referred to as non-minimal coupling) [15–18].

Here we address the question of whether there are any classes of potentials, besides the exponential, yielding the self-similar behaviour found in the latter case. In general the motivation for considering exponential potentials is drawn from supergravity (superstrings) where these type of potentials commonly arise after the dimensional reduction into a 4-dimensional space-time theory. However the resulting effective potentials are seldom simply exponentials and exhibit some form factors due to couplings between fields. On other hand, it has been claimed that the only potentials inducing the tracker scaling solutions are the exponential potentials [9]. In the present work we perform a phase-plane analysis of the complete dynamical system corresponding to a flat FRW cosmological models with a perfect fluid and a self-interacting scalar field and show that every positive and monotonous potential which is asymptotically exponential also yields a tracking solution. Our results are the outcome of taking into consideration the full third order dynamical system. In previous work in the literature the analysis which were carried out miss an equation for the scalar field and the resulting reduced system is only appropriate to study the exponential potential cases.

## II. SCALAR FIELD COSMOLOGIES

Consider the flat homogeneous and isotropic universes given by the Friedmann-Robertson-Walker (FRW) metric

$$ds^2 = -dt^2 + a^2(t) \left[ dr^2 + r^2(d\theta^2 + \sin^2 \theta d\phi^2) \right]. \quad (1)$$

Assume that the matter sources are a perfect fluid characterized by the equation of state  $p = (\gamma - 1)\rho$ , where  $0 \leq \gamma \leq 2$  is a constant, and a self-interacting scalar field  $\varphi$  with the potential  $V(\varphi)$ . The field equations then read

$$H^2 = \frac{8\pi G}{3} \left( \frac{\dot{\varphi}^2}{2} + V(\varphi) + \rho \right), \quad (2)$$

$$\ddot{\varphi} + 3\frac{\dot{a}}{a}\dot{\varphi} = -\frac{\partial V(\varphi)}{\partial \varphi}, \quad (3)$$

where the overdots stand for the derivatives with respect to the time  $t$ , and  $H = \dot{a}/a$ . Another equation which is useful is

$$\dot{H} = -\frac{\aleph}{2} (\dot{\varphi}^2 + \gamma\rho). \quad (4)$$

Following [8,12] we introduce the new time variable  $N = \ln \tilde{a}$  and the expansion normalized variables

$$x = \frac{\aleph\dot{\varphi}}{\sqrt{6}H} \quad (5)$$

$$y = \frac{\aleph V(\varphi)}{\sqrt{3}H} \quad (6)$$

where  $\aleph^2 = 8\pi G$  (in what follows we shall adopt units that set  $\aleph^2 = 1$ ).

We thus obtain the following third order dynamical system

$$(x)' = -3x - \sqrt{\frac{3}{2}} \left( \frac{\partial_\varphi V}{V} \right) y^2 + \frac{3}{2}x [2x^2 + \gamma(1 - x^2 - y^2)] \quad (7)$$

$$y' = \sqrt{\frac{3}{2}} \left( \frac{\partial_\varphi V}{V} \right) xy + \frac{3}{2}y [2x^2 + \gamma(1 - x^2 - y^2)], \quad (8)$$

$$\varphi' = \sqrt{6}x. \quad (9)$$

where we have used  $\rho/3H^2 = \Omega_m = 1 - x^2 - y^2$ .

The latter equation (9) is usually overlooked and only the former two are considered in the previous work on scaling solutions. This is due to the fact that if one considers an exponential potential  $V(\varphi) \propto \exp(\lambda\varphi)$  the two equations, (7) and (8), form a two dimensional closed autonomous system of the polynomial type whose qualitative behaviour can be easily determined, and which gives rise to scaling (tracker) solutions [8,12].

Here, we shall continue to consider the full system. We immediately see from the consideration of equations (7-9) that they permit the study of any other positive potentials  $V(\varphi)$  besides the exponential. Whenever  $V(\varphi) \neq V_0 \exp(\lambda\varphi)$  equation (9) tells us that all singular points must have  $x = 0$ . The former equations then reduce to

$$x' = -\sqrt{\frac{3}{2}} \partial_\varphi \ln V y^2 \quad (10)$$

$$y' = \frac{3}{2} \gamma y(1 - y^2), \quad (11)$$

We may rule out the line of equilibria  $(x = 0, y = 0, \varphi)$ , which corresponds to a fluid dominated universe and is unstable in the  $y$  direction with eigenvalue  $3\gamma/2 \geq 0$ . Scaling solutions, in the sense of power law behaviour for the scalar field energy density, for fluid dominated universes were studied by Liddle and Scherrer [9]. In this approximation, they found a late time attractor for which both the kinetic and the potential energy of the

field behave as a certain power of the inverse scale factor, while the dominant component scales as a smaller power of the inverse scale factor, for potentials with null extrema. Their analysis cannot be retrieved through the study of system (7-9), which becomes singular when  $V(\varphi) = 0$ . Anyway, our interest is the search for trackers, that is, solutions for which the field's energy density (and its kinetic energy and potential as well) and the fluid's energy density all scale as the same power of the scale factor. And this means equilibria of (7-9) in the region  $0 < x^2 + y^2 < 1$ . Considering again the system (7-9), we see that there exists only another class of fixed points, namely those of the form

$$p_{crit} = (x = 0, y = 1, \varphi = \varphi_0), \quad (12)$$

where  $\varphi_0$  is such that  $\partial_\varphi V = 0$ . This equilibrium point is degenerate, but it is easy to check taking into account the second order terms in the  $y'$  equation that it is stable provided that  $\varphi_0$  is a minimum of  $V(\varphi)$ . Again, this is no tracker solution. In this case, we have the opposite scenario, a field dominated universe as a result of an effective cosmological constant.

This analysis has lead many authors to claim that only exponential potentials may give rise to tracker solutions [9]. However, we must still explore the possibility of the existence of equilibria at  $\varphi = \pm\infty$ . In fact, the phase space of the full three dimensional system (7-9) is not compact, and there may exist late time attractors with the required behaviour on the boundary  $\varphi = \pm\infty$ . In what follows, we shall focus on a particular potential,

$$V(\varphi) = \left( \frac{2\Gamma}{\alpha + \alpha c(\varphi)} \right)^{1/(\lambda-1)} \left( \frac{(2 + \Gamma) + (2 - \Gamma)c(\varphi)}{2(1 + c(\varphi))} \right), \quad (13)$$

where

$$c(\varphi) = \cosh(\sqrt{3\Gamma}(\lambda - 1)\varphi), \quad (14)$$

and  $\Gamma$ ,  $\alpha$  and  $\lambda > 1$  are positive constants. This particular potential is motivated by the fact that, in the absence of any other matter component, it corresponds to an equation of state of the form

$$p + \rho = \Gamma\rho - \alpha\rho^\lambda, \quad (15)$$

which is a natural generalization of the 'exponential' equation of state

$$p + \rho = \gamma\rho. \quad (16)$$

This is only to fix ideas, and it will become clear that the only essential feature of the potential is that it should behave asymptotically as an exponential potential. Barrow [20] has derived exact solutions of string-driven cosmologies introducing a bulk viscosity pressure which yields an equation of state equivalent to (15) when the spatial curvature vanishes.

Introducing the auxiliary variable  $\psi = 1/\varphi$ , we obtain

$$\psi' = -\psi^2\sqrt{6}x \quad (17)$$

which shows that the 'infinity manifold'  $\psi = 0$  is invariant. The flow on this invariant manifold is given by the following equations obtained from Eqs. (7,8) in the limit  $\psi \rightarrow 0$

$$(x)' = -3x + k\sqrt{\frac{3}{2}}y^2 + \frac{3}{2}x [2x^2 + \gamma(1 - x^2 - y^2)] \quad (18)$$

$$y' = -k\sqrt{\frac{3}{2}}xy + \frac{3}{2}y [2x^2 + \gamma(1 - x^2 - y^2)] , \quad (19)$$

where  $k = \sqrt{3\Gamma}$ . This system was studied in [8,12] (note that in Ref. [8] the potential is  $V = V_0 \exp -\lambda\varphi$ , i.e.,  $\lambda$  is used instead of  $k$ ), and it has a sink whenever  $k^2 > 3\gamma$  located at  $(x = \sqrt{3/2\gamma}/k, y = \sqrt{3(2-\gamma)\gamma/(2k^2)})$ . Since the  $x$  coordinate for this sink is positive, equation (17) tells us that this equilibrium point is a sink for the full regularized system on the  $\varphi = +\infty$  invariant boundary. Moreover, this equilibrium point is the only stable equilibrium of the full system, and will therefore attract almost all the initial conditions. Therefore, just as in the case of the exponential potential, the only late time attractor for the system for these values of the parameters is a tracker scaling solution. This conclusion carries over to every positive and monotonous potential, provided that its asymptotic behaviour is exponential. In fact for all these potentials the projection of the late-time flow on the  $(x, y)$ -plane is qualitatively similar to the flow of the autonomous system (18,19) associated with the exponential potential.

### III. DISCUSSION

The main conclusion of our work is that tracker solutions are compatible with a wide class of potentials and thus are not restricted to the exponential potential case. The potentials must be positive and monotonous, and have to asymptotically approach the exponential behaviour. As the qualitative behaviour of the models based on the extended class of tracker potentials discussed in this paper is basically the same as that of the models based on exponential potentials, they will face the same sort of difficulties as those faced by the latter [8,7,11,9] in the later stages of the evolution of the universe.

However in the early universe the differences that are allowed with regard to the exponential potential permit greater flexibility in building models with scalar fields. For instance the self-interacting scalar field cosmologies based on the potential of Eq. (13) have an early-time regime dominated by the scalar field (when  $x, y \simeq 0$ ). During this stage the effective equation of state is  $\gamma_\varphi \simeq \Gamma s/(s + \alpha)$ , where  $s = (a/a_0)^{3(\lambda-1)\Gamma}$  and thus, for  $s$  sufficiently small,  $\gamma_\varphi \simeq 0$  yielding slow-roll inflation. Indeed the relevant part of the potential during such a stage is  $V(\varphi) \simeq (\Gamma/\alpha)^{1/(\lambda-1)}(1 - \nu\varphi^2)$ , where  $\nu = 3\Gamma(\lambda - 1)[\lambda - (2 - \Gamma)(\lambda - 1)]/8$  is a constant, and is akin to the pseudo-Nambu-Goldstone boson of natural inflation [21] which gives slow-roll without fine-tuning [22]. The required minimum number of e-foldings is guaranteed by  $3\Gamma(\lambda - 1) \lesssim \ln \alpha/60$ , and hence does not restrict the value of  $\Gamma$ . Since the potential (13) eventually yields the tracker behaviour (when  $\Gamma > \gamma_{matter}$ ), it permits a smooth transition out of quasi-de Sitter inflation into an epoch dominated by the matter. In particular, if  $\Gamma > \gamma_{matter} = 4/3$ , the scalar field self-adjusts to the radiation behaviour and the scale factor evolves as a power-law  $a(t) \propto t^{1/2}$ . Such a smooth transition cannot be achieved with a bare exponential potential which either yields eternal inflation (whenever  $k^2 < 3\gamma_{matter}$ ) or, alternatively, yields the tracking behaviour without any stage of inflation. In this latter case inflation is usually fulfilled through the additional consideration of yet

another scalar field devoted to that purpose. The simpler model based on the potential (13) thus illustrates the underlying possibilities of the wider class of tracking potentials presented in this work to try and meet the various theoretical and observational constraints that have to be faced by a scalar field cosmology.

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